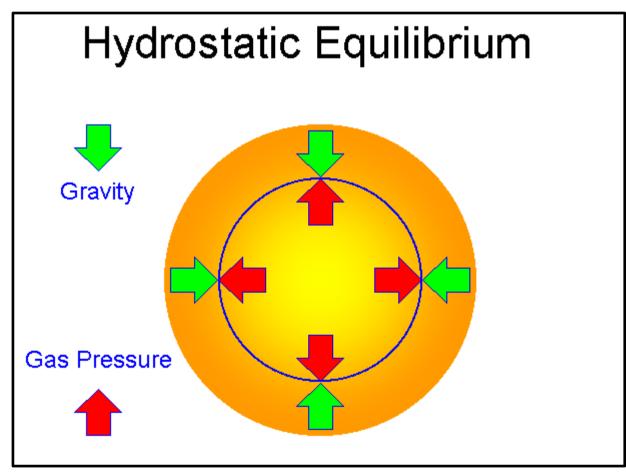


Secuencia principal

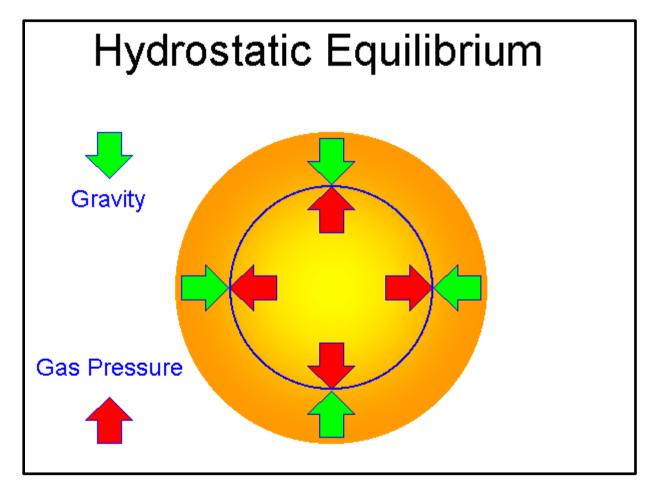
Una estrella esta en secuencia principal cuando logra el equilibrio hisdrostático.

Secuencia principal



principal cuando

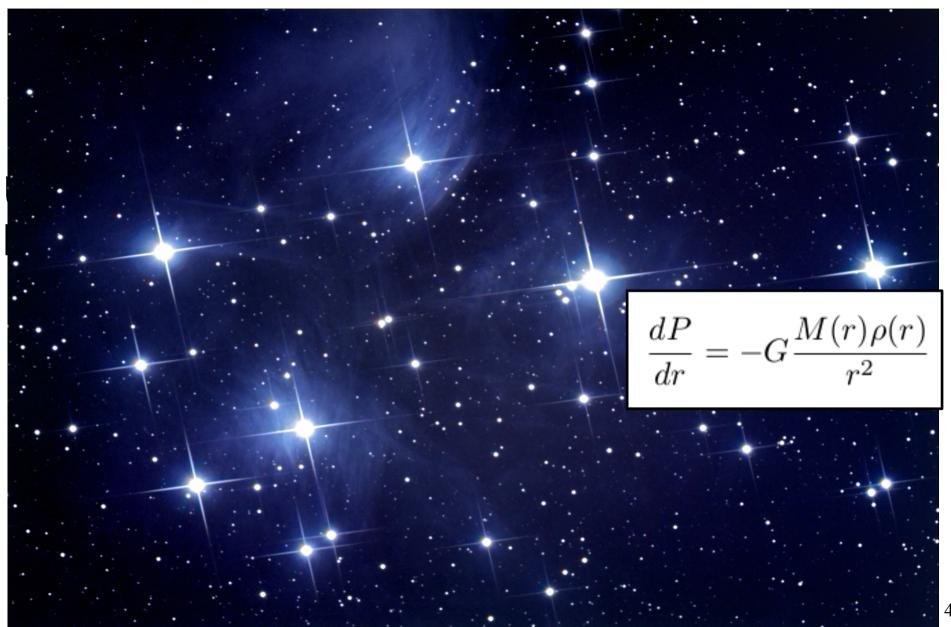
Secuencia principal



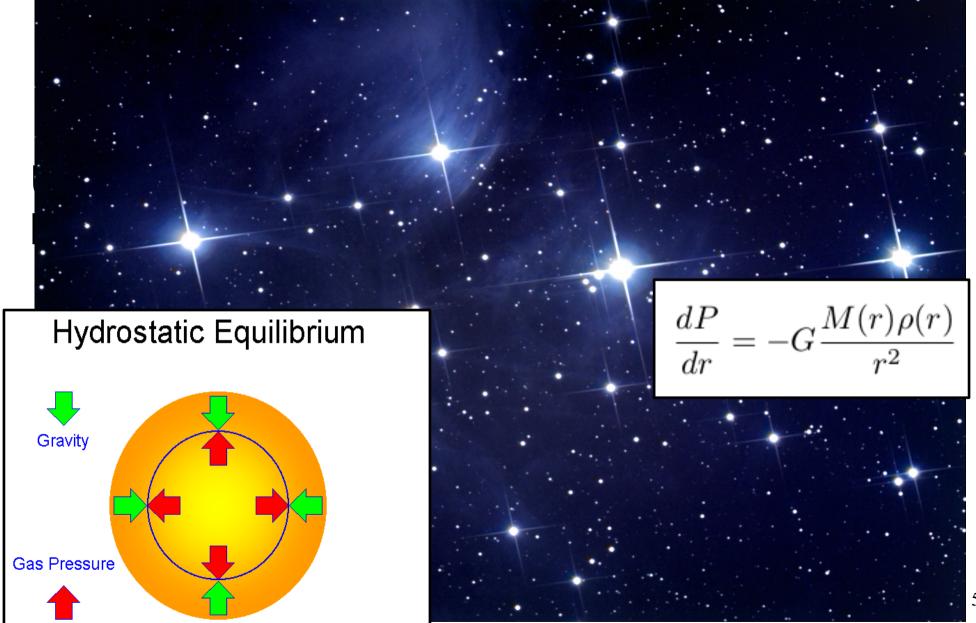
principal cuando

$$\frac{dP}{dr} = -G\frac{M(r)\rho(r)}{r^2}$$





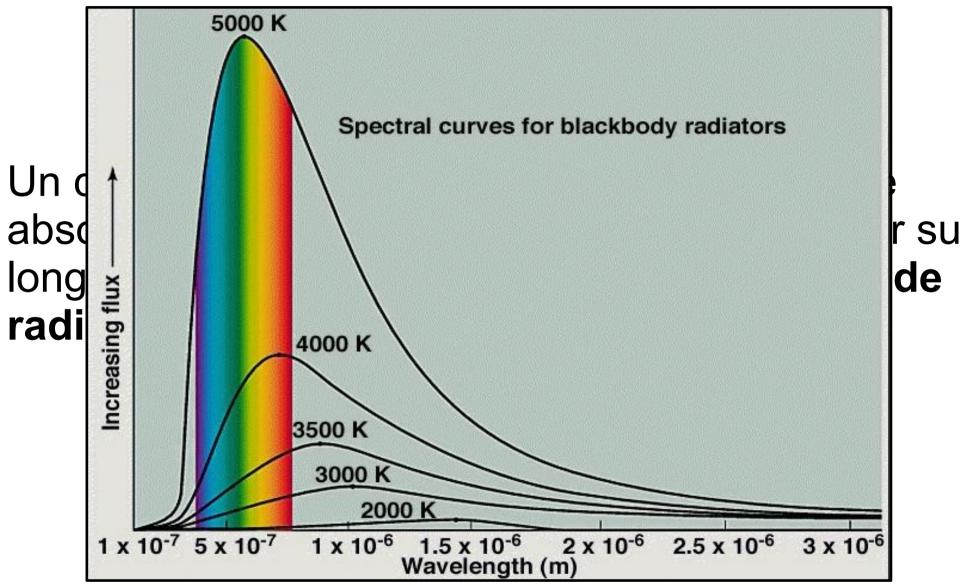




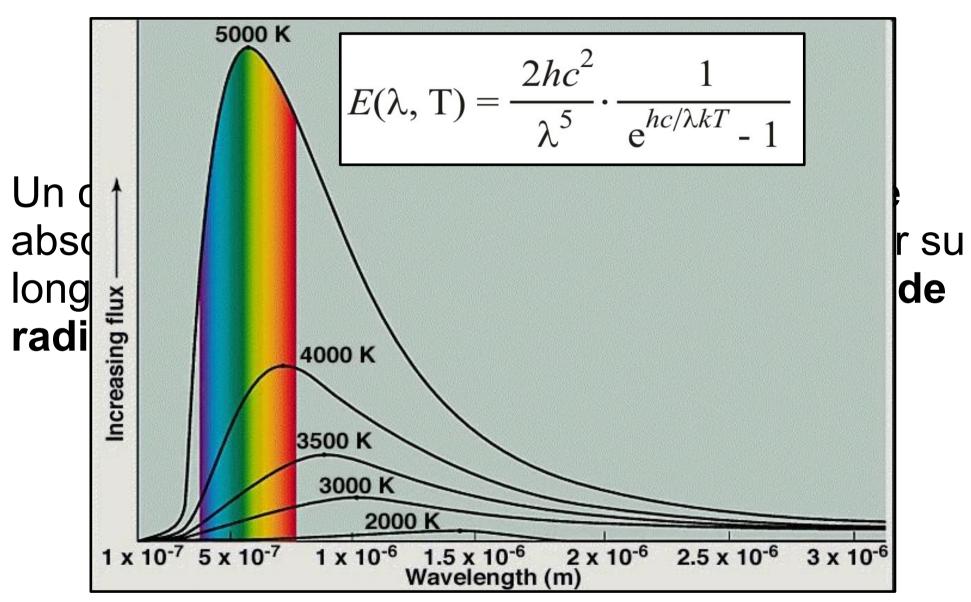
Radiación de cuerpo negro

Un cuerpo negro es un sistema físico ideal que absorbe toda la radiación incidente sin importar su longitud de onda: es un absorbente perfecto de radiación electromagnética

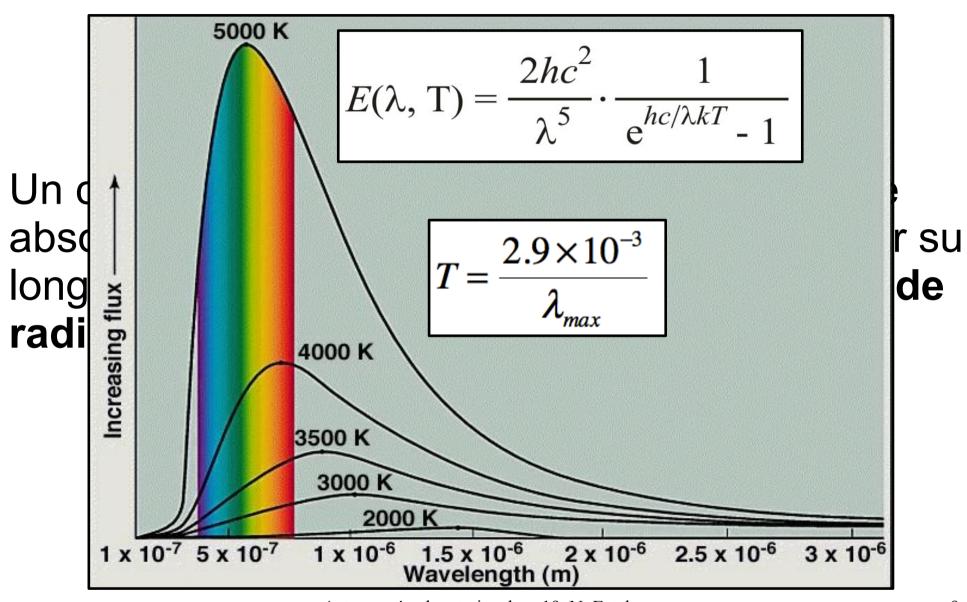




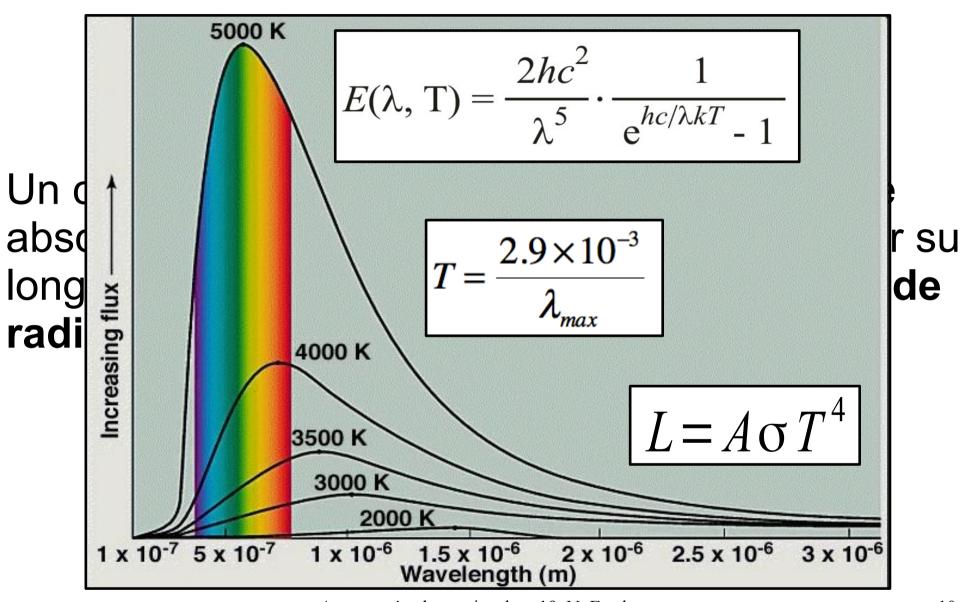
Astronomía planetaria, clase 19. N. Estelar



Astronomía planetaria, clase 19. N. Estelar

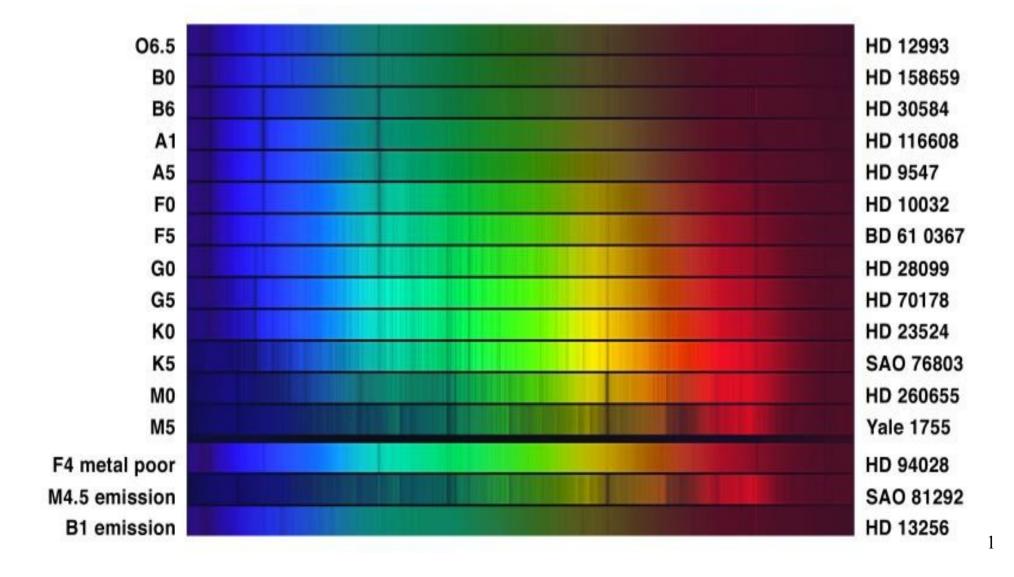


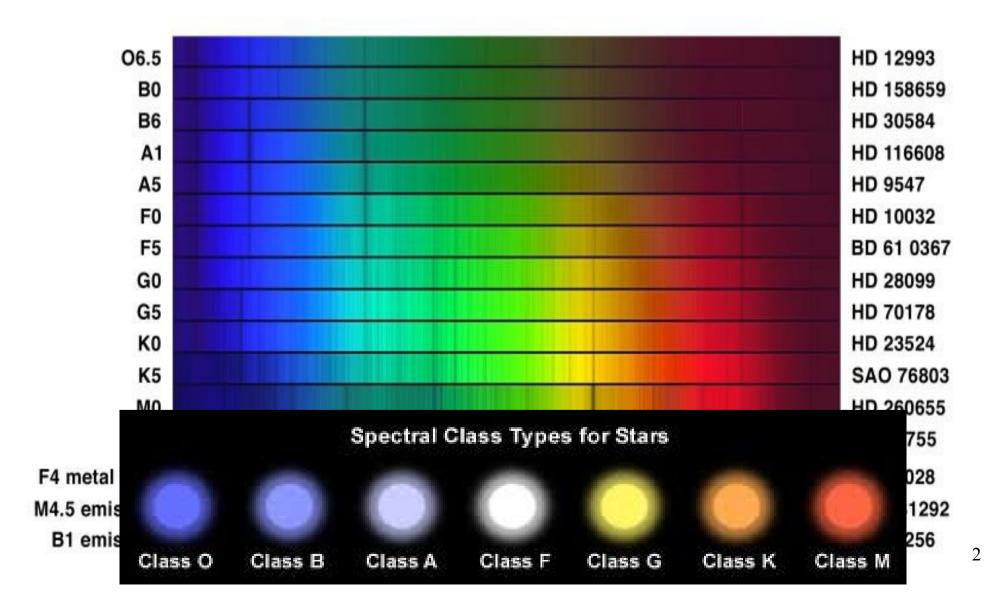
Astronomía planetaria, clase 19. N. Estelar



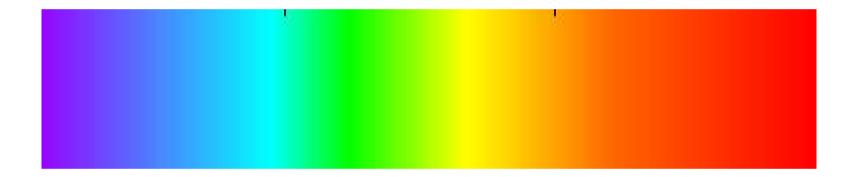
Astronomía planetaria, clase 19. N. Estelar





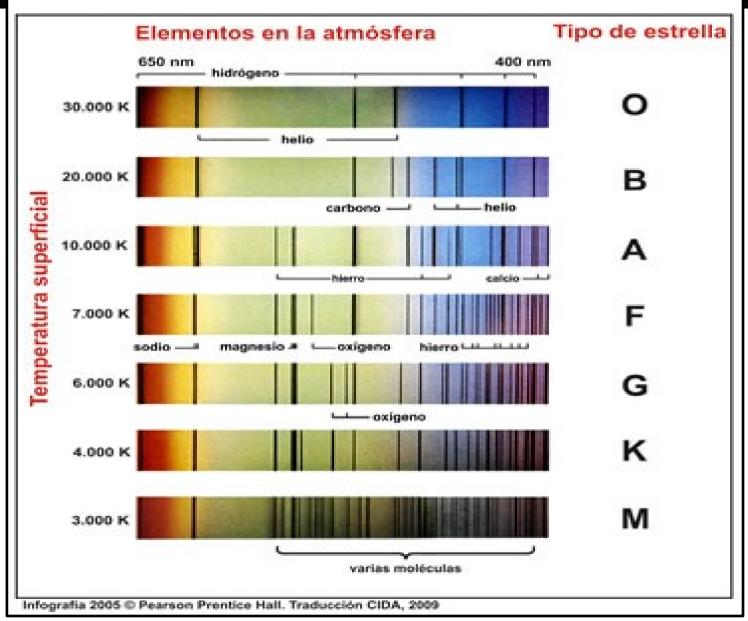


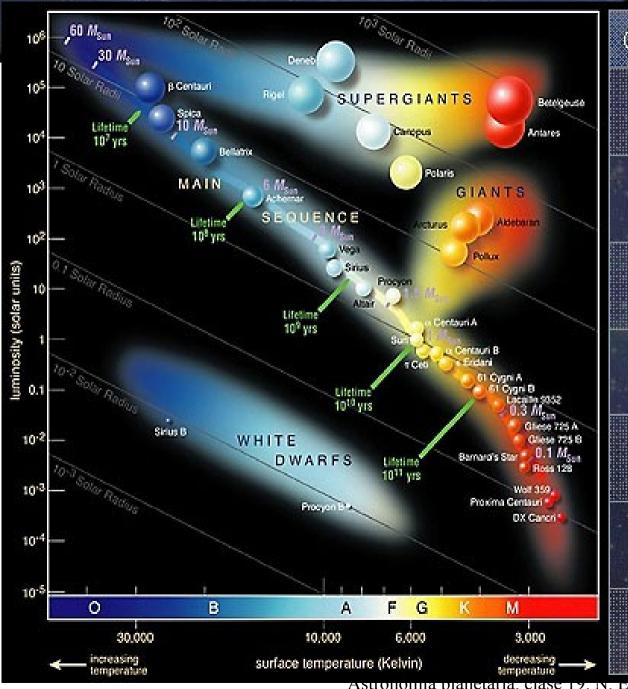




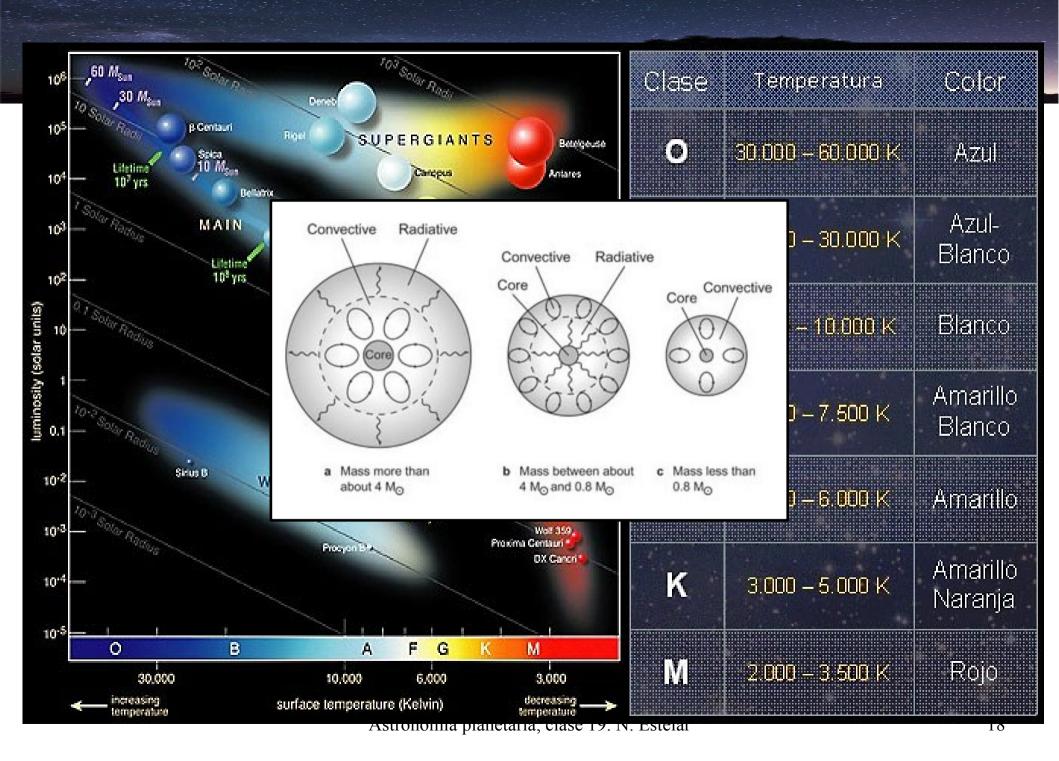
Class	Color	Prominent Spectral Lines	Surface Temp. (K)
0	Blue	Ionized helium, hydrogen	> 25,000 K
В	Blue-white	Neutral helium, hydrogen	11,000 – 25,000 K
Α	White	Hydrogen, ionized sodium and calcium	7,500 – 11,000 K
F	White	Hydrogen, ionized and neutral sodium and calcium	6,000 – 7,500 K
Ð	Yellow	Neutral sodium and calcium, ionized calcium, iron, magnesium	5,000 – 6,000 K
Κ	Orange	Neutral calcium, iron, magnesium	3,500 – 5,000 K
М	Red	Neutral iron, magnesium, and neutral titanium oxide	< 3,500 K

CLASIFICACIÓN ESPECTRAL DE LAS ESTRELLAS

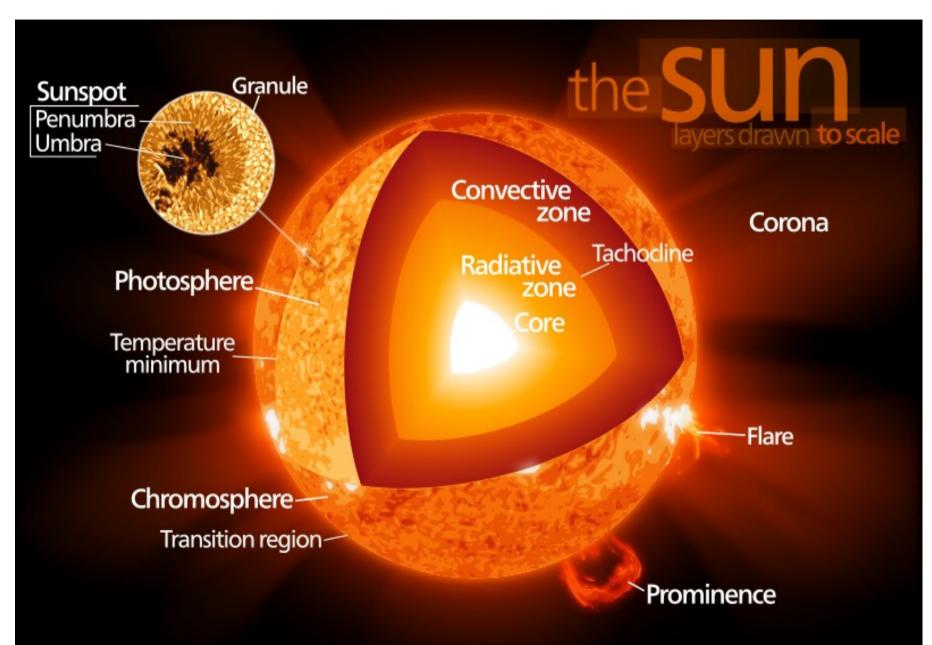


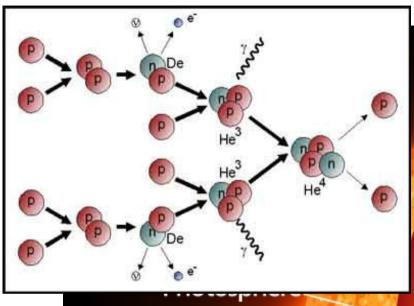


Clase	Temperatura	Color
0	30.000 – 60.000 K	Azul
B	10.000 – 30.000 K	Azul- Blanco
A	7.500 – 10.000 K	Blanco
F	6.000 – 7.500 K	Amarillo Blanco
G	5.000 – 6.000 K	Amarillo
K	3.000 – 5.000 K	Amarillo Naranja
M	2(80)) – 6,500 K	Rojo



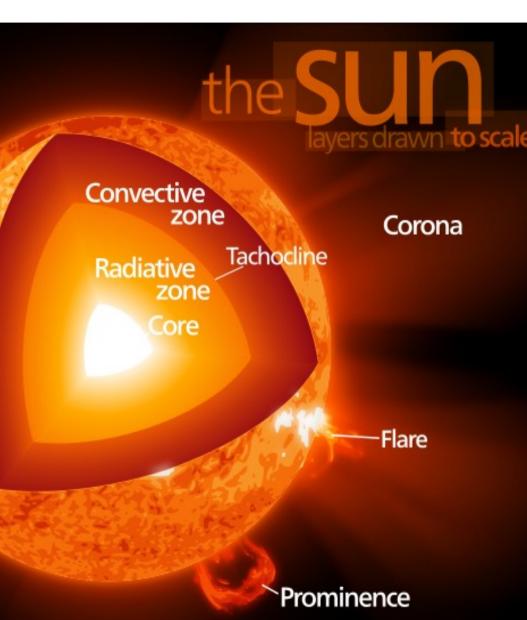
Estrellas de baja masa (Sol)

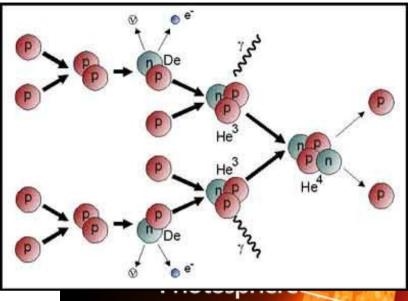




Temperature minimum

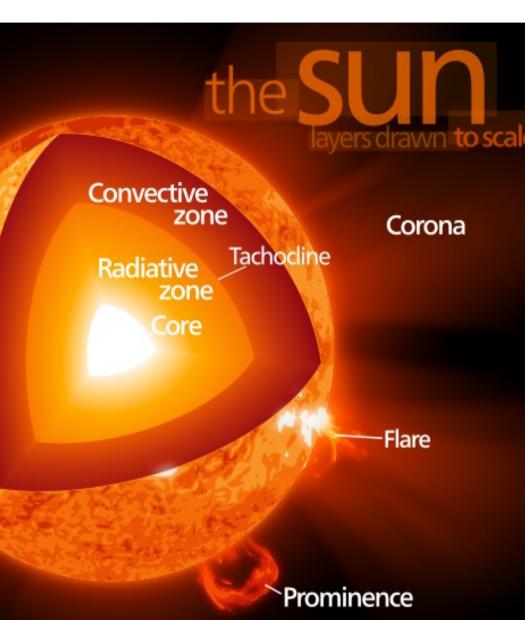
Chromosphere Transition region

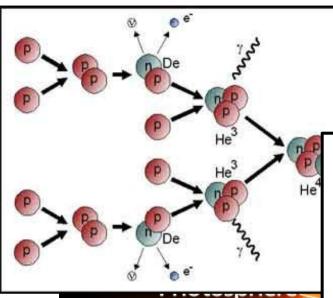




Temperature minimum

Chromosphere Transition region





 ${}_{1}^{1}\text{H} + {}_{1}^{1}\text{H} \rightarrow {}_{1}^{2}\text{H} + e^{+} + \nu_{e}$

 $_{1}^{2}\text{H} + _{1}^{1}\text{H} \rightarrow _{2}^{3}\text{He} + \gamma$ 69% 31%

99.7%

 ${}_{2}^{3}\text{He} + {}_{2}^{3}\text{He} \rightarrow {}_{2}^{4}\text{He} + 2 {}_{1}^{1}\text{H}$

 ${}_{2}^{3}\text{He} + {}_{2}^{4}\text{He} \rightarrow {}_{4}^{7}\text{Be} + \gamma$

(PPI)

 $^{7}_{4}$ Be + $e^{-} \rightarrow ^{7}_{3}$ Li + ν_{e}

 $_{3}^{7}\text{Li} + _{1}^{1}\text{H} \rightarrow 2_{2}^{4}\text{He}$

(PP II)

 ${}^{7}_{4}\mathrm{Be} + {}^{1}_{1}\mathrm{H} \rightarrow {}^{8}_{5}\mathrm{B} + \gamma$

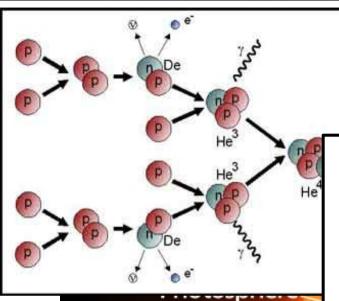
0.3%

 ${}_{5}^{8}\mathrm{B} \rightarrow {}_{4}^{8}\mathrm{Be} + e^{+} + \nu_{e}$

 $^{8}_{4}\text{Be} \rightarrow 2\,^{4}_{2}\text{He}$ (PP III)

Temperature minimum

Chromospher Transition re



the SUN

 ${}_{1}^{1}H + {}_{1}^{1}H \rightarrow {}_{1}^{2}H + e^{+} + \nu_{e}$

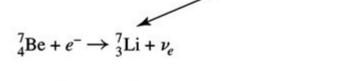
99.7%

 $_{1}^{2}\text{H} + _{1}^{1}\text{H} \rightarrow _{2}^{3}\text{He} + \gamma$

26.73 MeV - ^{7/3 Be + γ}

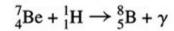
Temperature minimum

> Chromospher Transition re



$$_{3}^{7}\text{Li} + _{1}^{1}\text{H} \rightarrow 2_{2}^{4}\text{He}$$
(PP II)

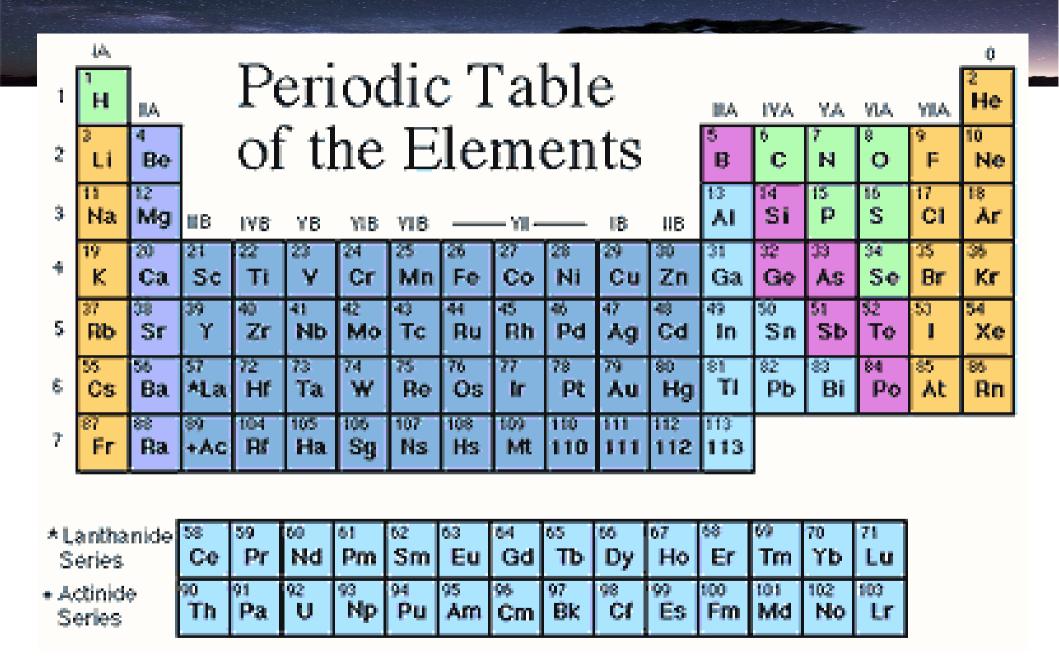
(PPI)

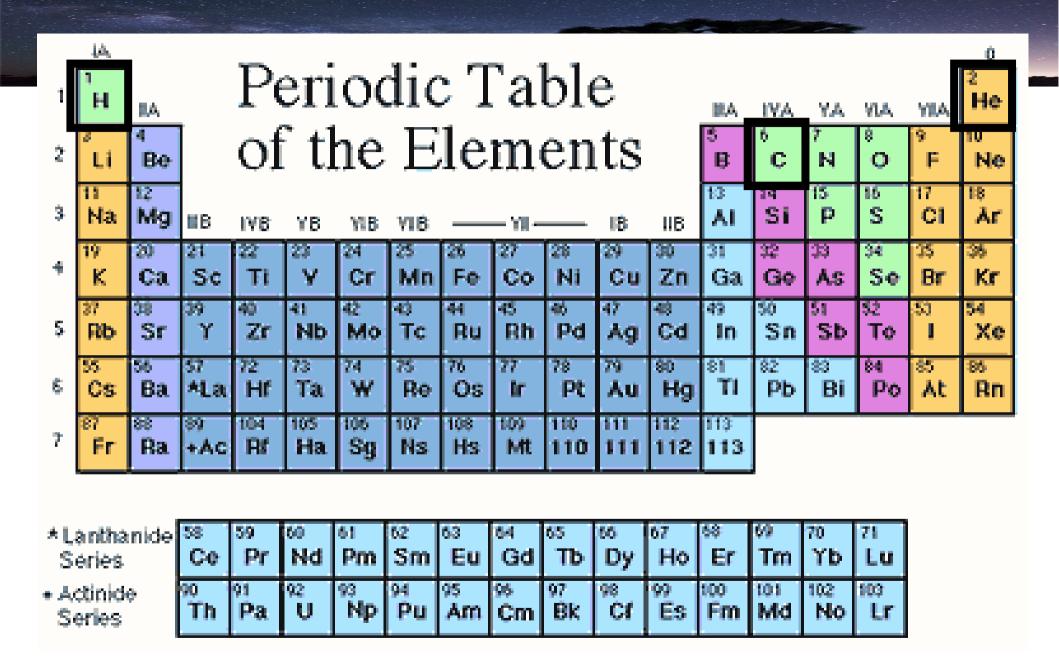


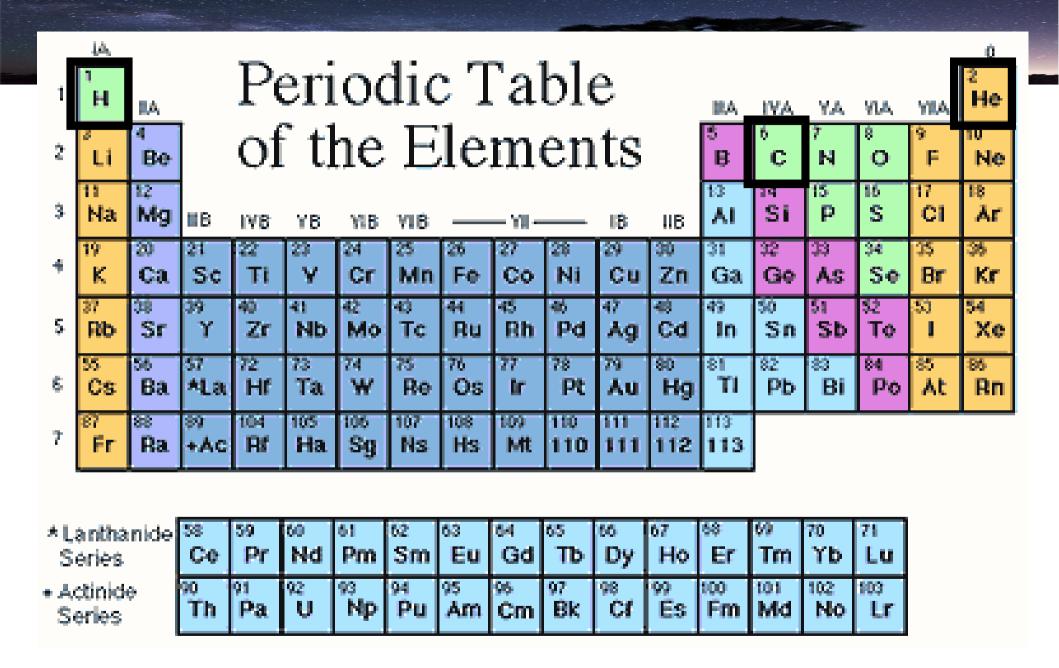
0.3%

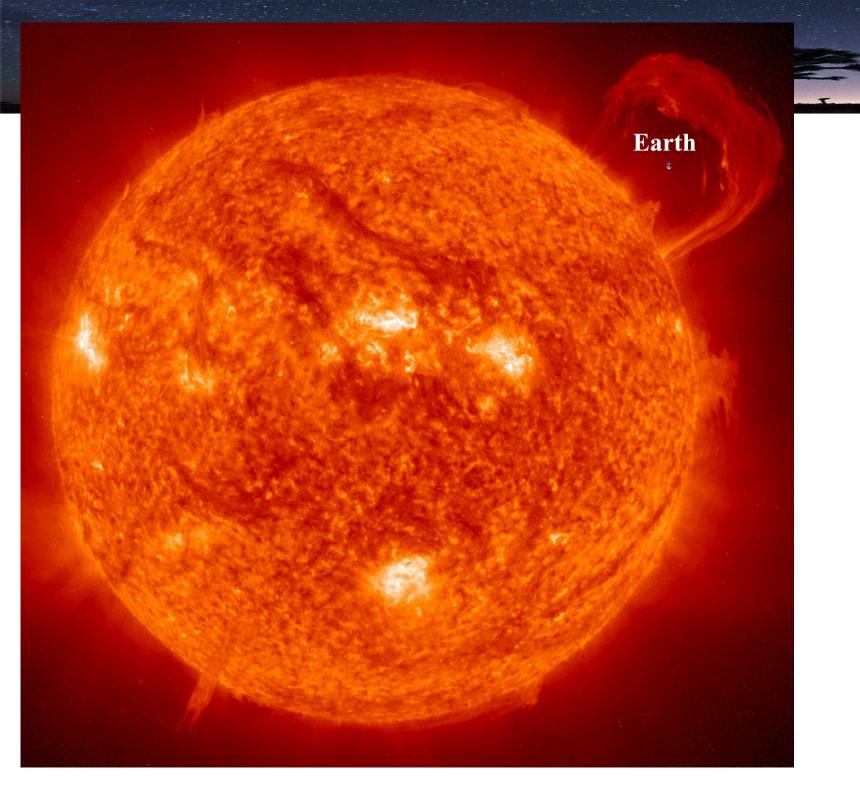
$${}^{8}_{5}\mathrm{B} \rightarrow {}^{8}_{4}\mathrm{Be} + e^{+} + \nu_{e}$$

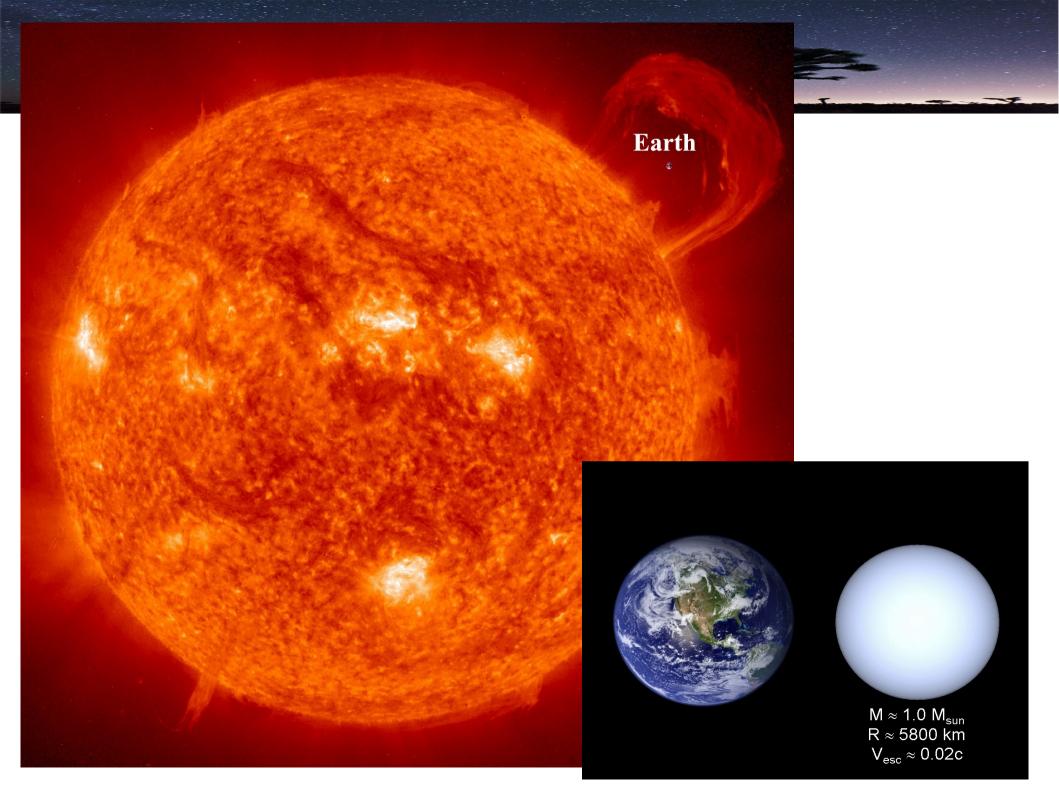
$${}_{4}^{8}\text{Be} \rightarrow 2 {}_{2}^{4}\text{He}$$
(PP III)

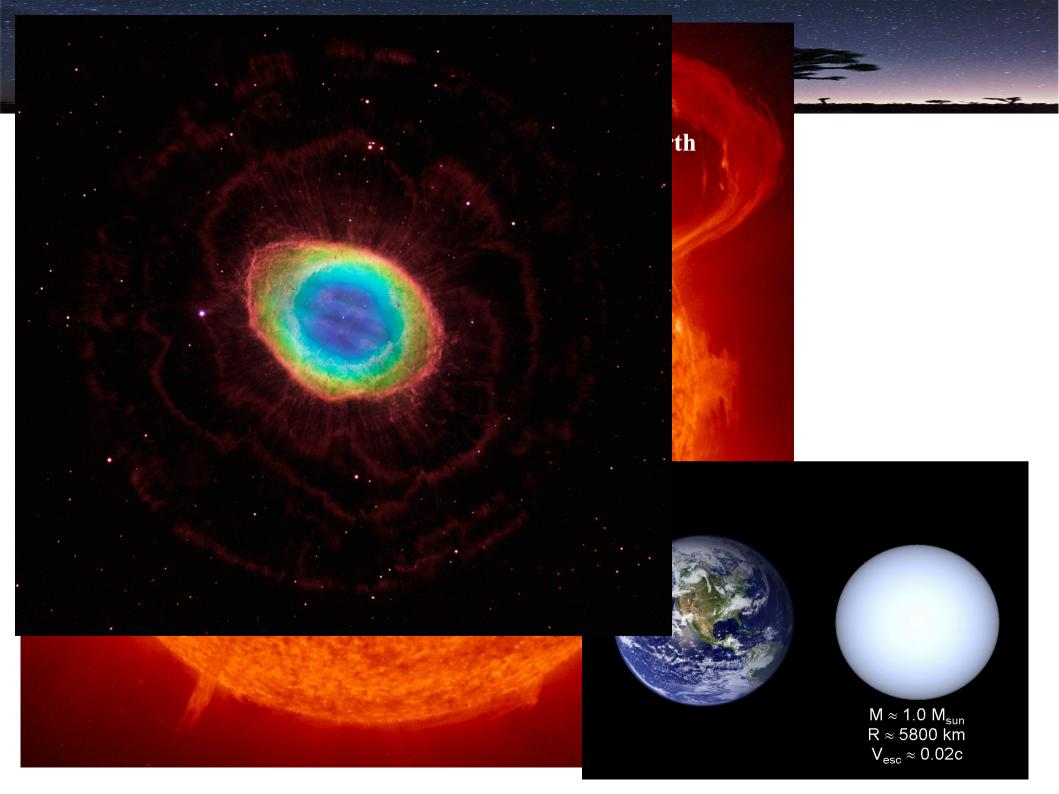


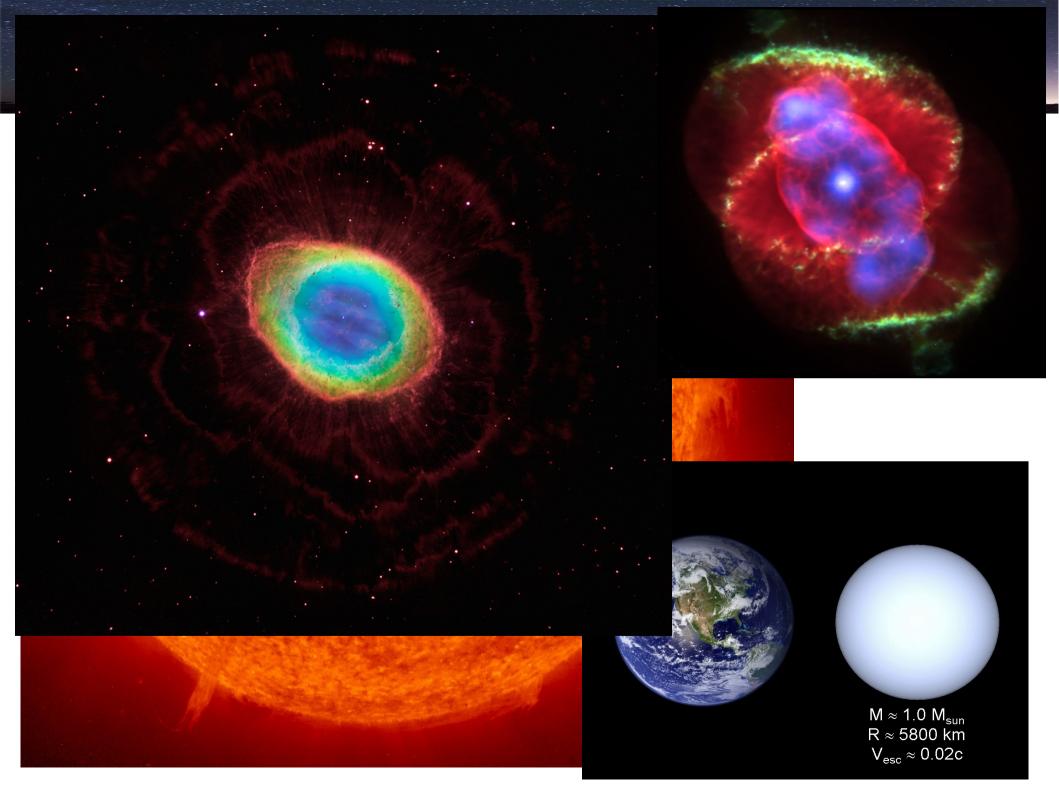


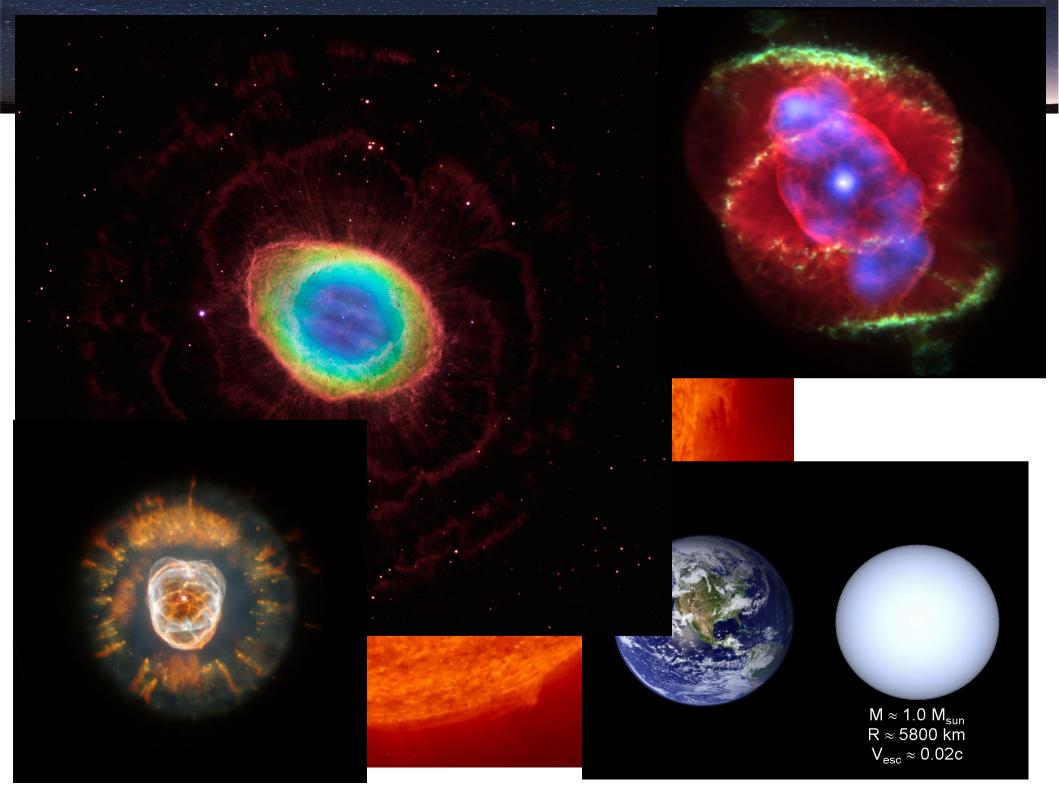


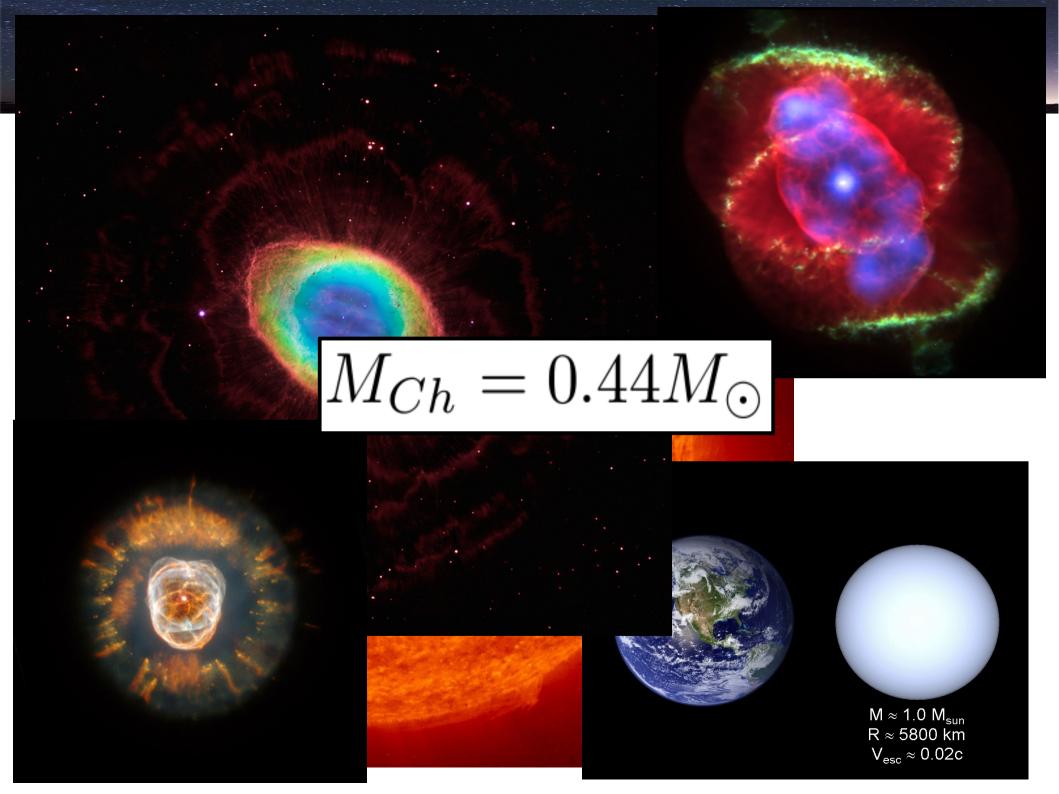






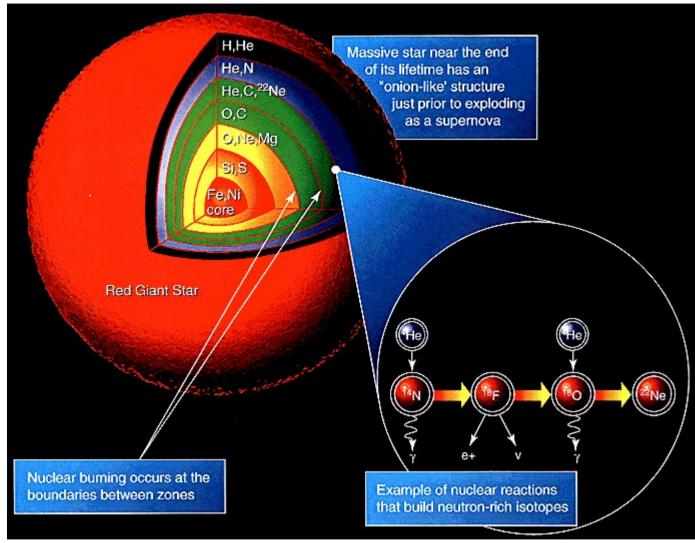






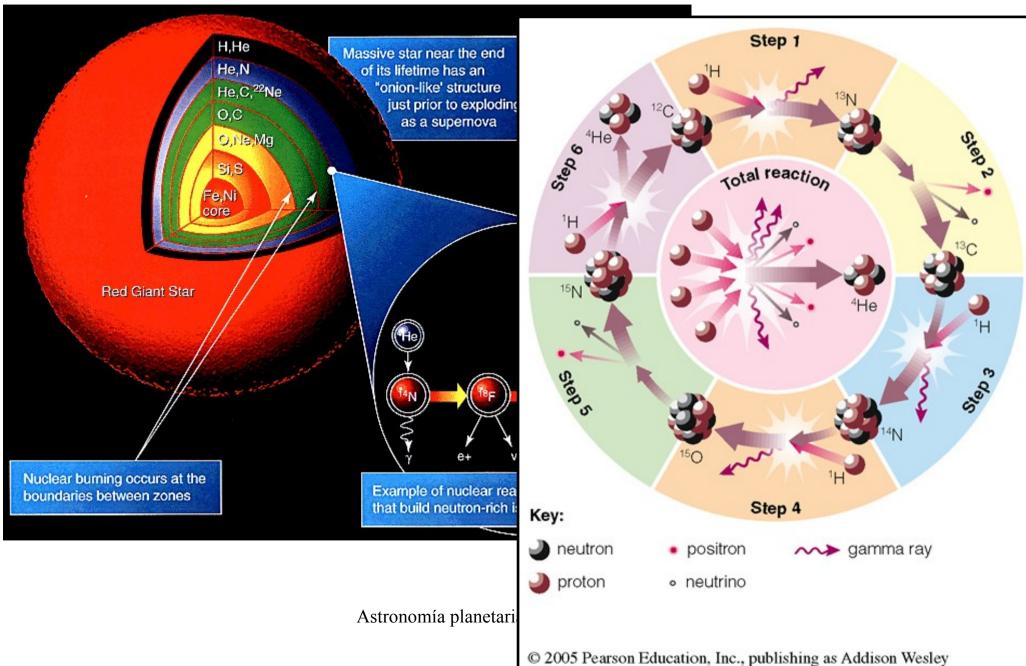
Estrellas de Alta masa (>10M_{sol})





(>10M_{sol})





Estrellas de Alta masa (>10M_{sol})

